



## **M33 with JAST/T80**



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CEFCOA  
LoRCA, kick-off workshop 2016 - Granada, Spain**



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and the J-PLUS Team**



# 1. Introduction: M33 & J-PLUS



## ◉ J-PLUS Survey:

- ◉ A large photometric survey at Observatorio Astrofísico de Javalambre in Teruel, Spain
- ◉ Filter Set: 12 Broad-, intermediate- and narrow-band filters (330 - 1000 nm)
- ◉ IFU-like capabilities

## Goals

- ◉ Test the capabilities of T80, T80Cam and the filter set
- ◉ Perform a detailed photometric analysis of the resolved and unresolved stellar population of M33
  - ◉ Unresolved Stellar Population:
    - ◉ A 2-D analysis of the underlying population
  - ◉ Star Cluster System
  - ◉ M33-M31 Interaction



# 1. Introduction: M33 & J-PLUS

## ● Hierarchical merging scenario:

Large disk galaxies derive from the merger and accretion of smaller subsystems

Where do dwarf spiral galaxies fit in the galaxy formation scenario?



## ● Why M33?

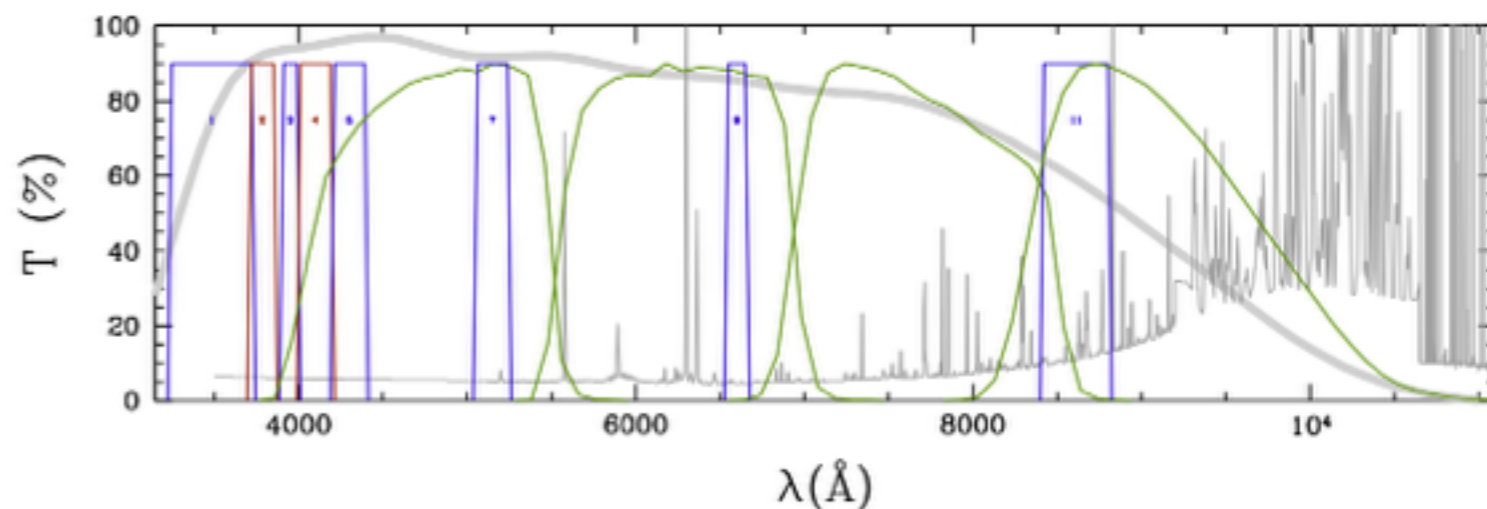
Distance = 870 kpc

The only nearby late-type spiral galaxy

Total mass ~ 20 times lower than the MW and only ~2 time greater than the Magellanic Clouds

Face-on view: study of its disk and halo populations

M33 provides a unique opportunity to test the  $\Lambda$ -CDM framework predictions in a regime different from the MW



Filter Number	Name	Central Wavelength (nm)	FWHM (nm)	Comments
JPLUS-01	u <sub>J</sub>	348.5	50.8	In common with J-PAS
JPLUS-02	F378	378.5	16.8	[OII]; in common with J-PAS
JPLUS-03	F395	395.0	10.0	Ca H+K
JPLUS-04	F410	410.0	20.0	H $\delta$
JPLUS-05	F430	430.0	20.0	G-band
JPLUS-06	g'	480.3	140.9	SDSS
JPLUS-07	F515	515.0	20.0	Mgb Triplet
JPLUS-08	r'	625.4	138.8	SDSS
JPLUS-09	F660	660.0	13.8	H $\alpha$ ; in common with J-PAS
JPLUS-10	i'	766.8	153.5	SDSS
JPLUS-11	F861	861.0	40.0	Ca Triplet
JPLUS-12	z'	911.4	140.9	SDSS



## 2. A Two-Dimensional Analysis

- ⊙ M33: Extended Object
  - ⊙ Unresolved Component

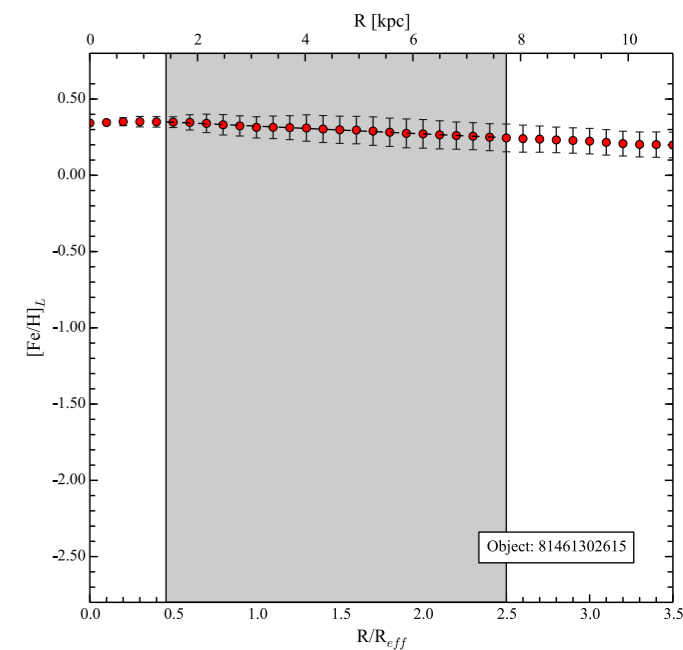
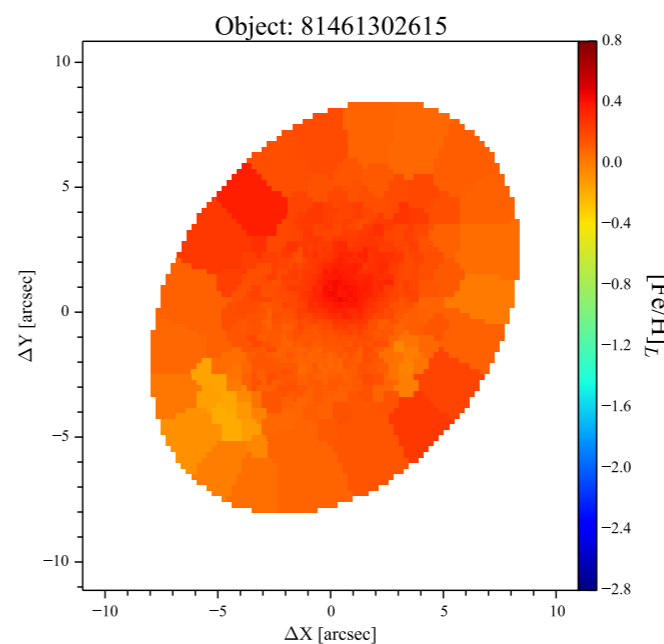
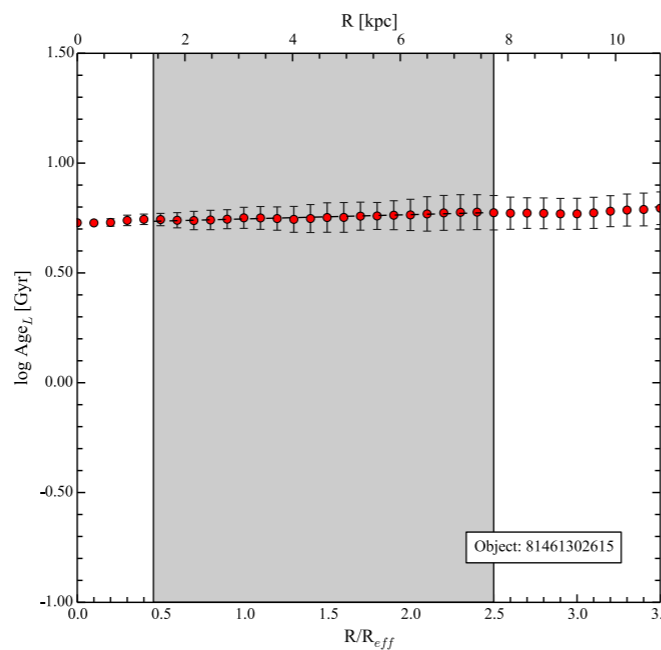
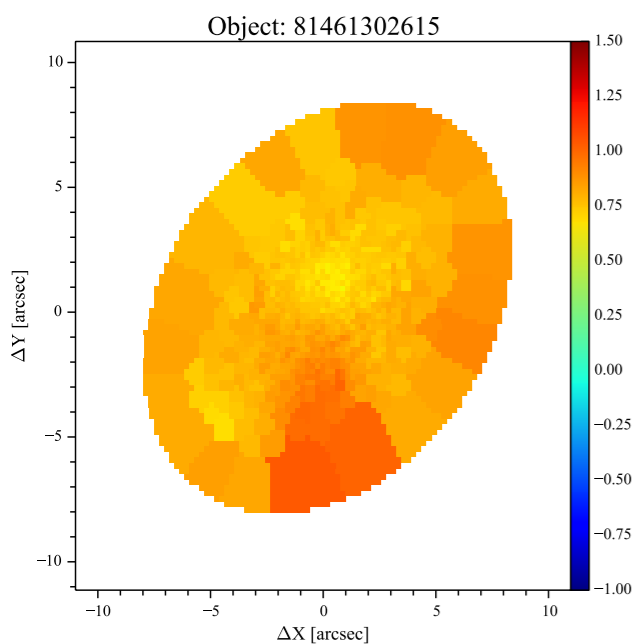


2-Dimensional Analysis



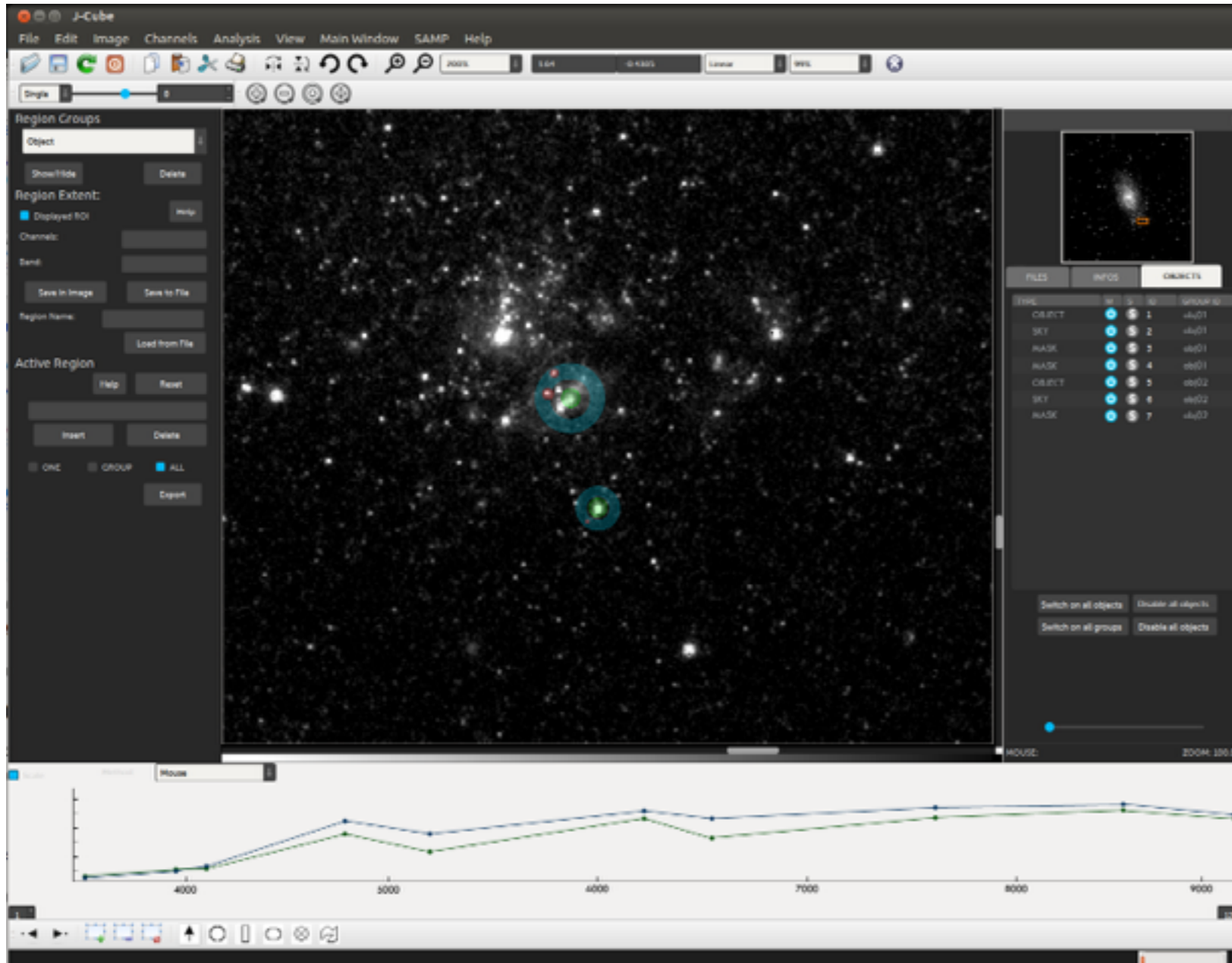
Spectral Fitting

- ⊙ Optimal set of filters: provides an adequate sampling of the galaxy SEDs.
- ⊙ Uniform and non-biased spatial sampling: allowing environmental studies
- ⊙ IFU-like capabilities: pixel-by-pixel investigation of the extended areas
- ⊙ Large survey area: a diversity of galactic components in a single telescope pointing





## 3. M33 Star Cluster System



- Distance: 870 kpc ( $m_v - M_v = 24.69$ )  
Reaching the faint tail of the luminosity function
- Multi-filter approach: detect star clusters and characterize their stellar population
- Ages and Metallicities for a big sample extending the most updated catalog of M33 star clusters (San Roman et al. 2010)

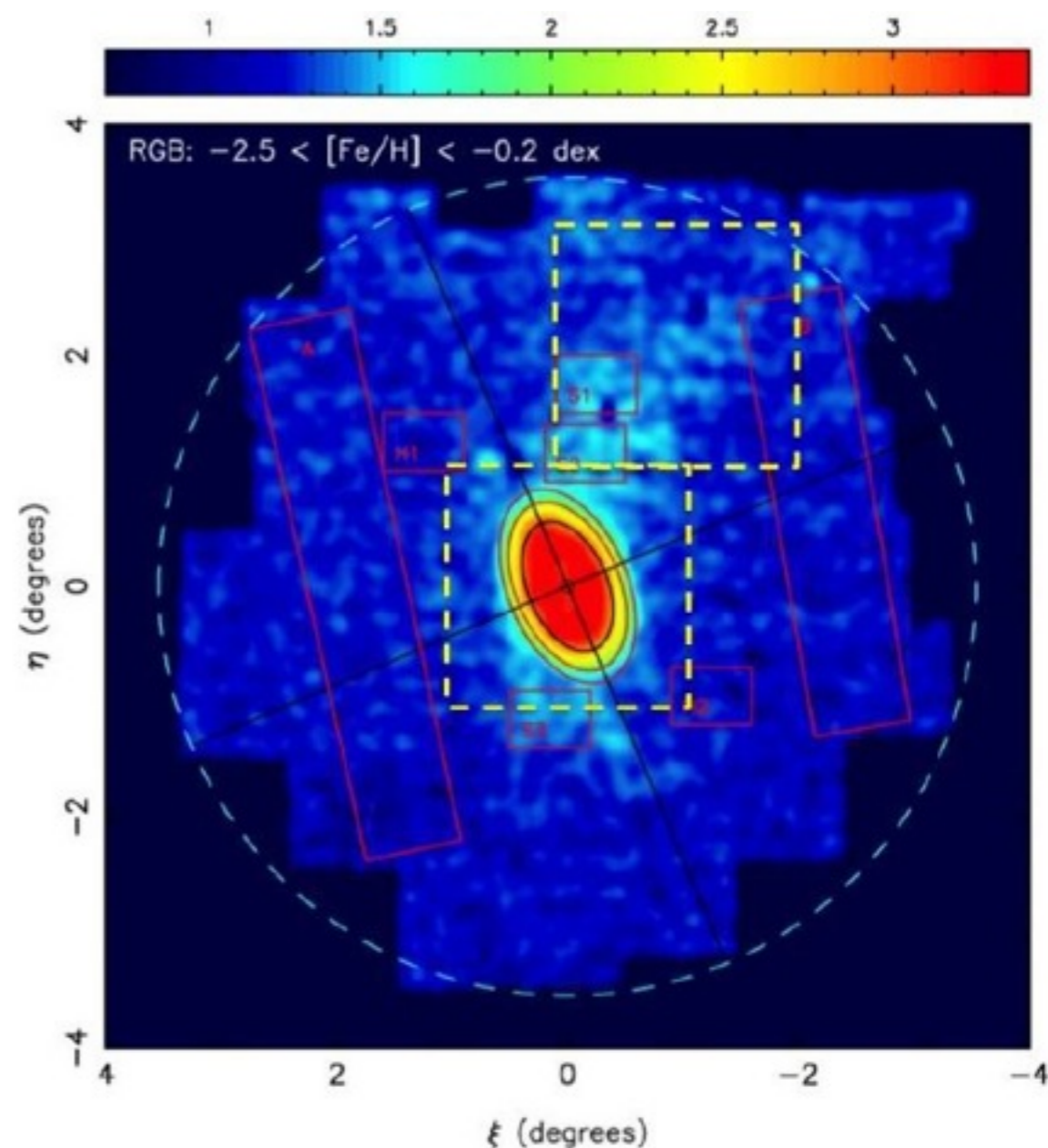


Perform a statistically solid inter comparisons of MW, M31 and M33



## 4. M33-M31 Interaction

- **McConnachie et al. 2009** find a previously unknown prominent stellar structure surrounding M33.
- An extension stretching  $\sim 2^\circ$  ( $\sim 30$  kpc; projected) to the northwest, towards M31
  
- **Two regions (3.9 squared degrees):**  
**M33-C:** centered on M33 and covering the disk and the outskirts  
**M33-S:** covering the line connecting M33-M31 to map the stellar substructure

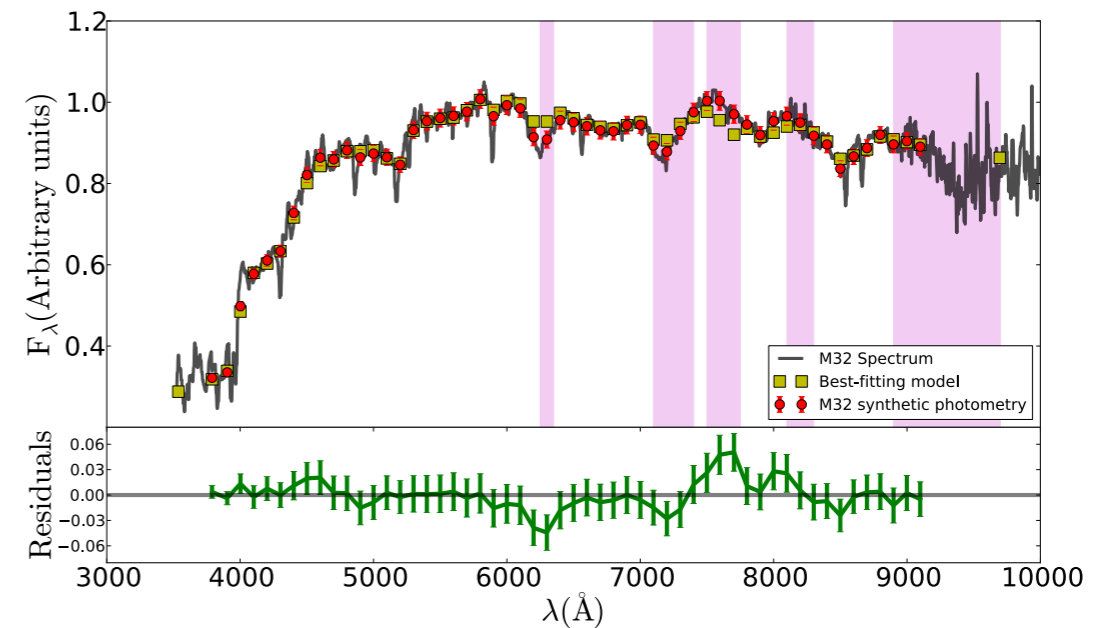


*McConnachie et al. 2010*



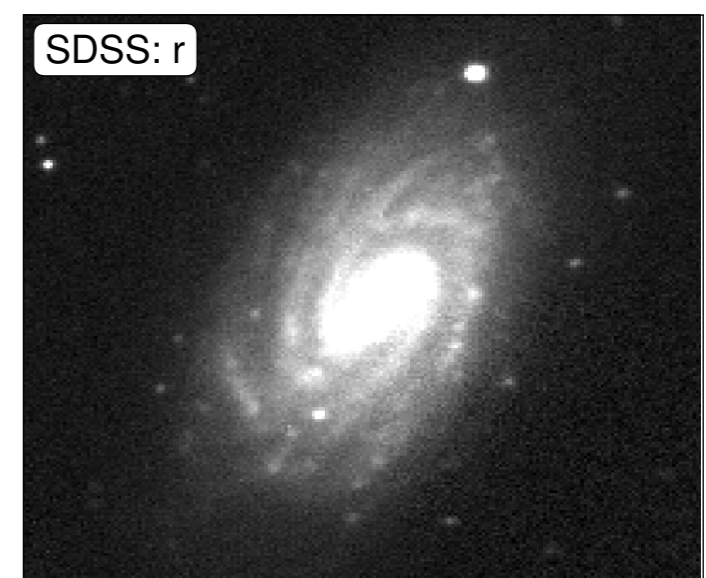
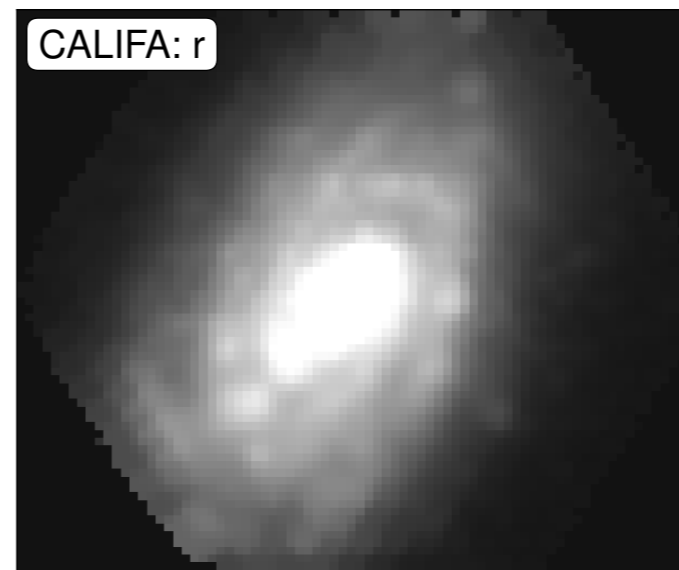
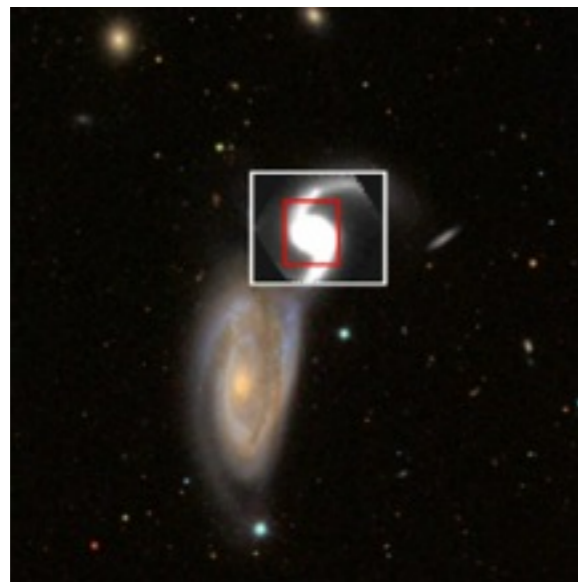
# 5. Scientific Potential

- 2D study of stellar population
  - J-PAS will permit the 2D study of stellar population in a very large sample of nearby galaxies
  - J-PAS photometry is the perfect anchor for the flux calibration of integral field surveys



- Spatial Resolution
  - Expected resolution of J-PAS:  $\sim 1''$  : smaller structures than local integral field surveys

- Environment Effect
  - Contiguous coverage of the northern sky: detailed studies of environmental effects





## 5. Summary

- ⦿ **Spectral fitting diagnostics of the resolved and unresolved populations will allow us to determine ages, metallicities and masses of the galactic disk, spheroidal component and cluster system:**
  - ⦿ **Are there any indications of substructure in the disk and halo of M33?**
  - ⦿ **Where is the disk/halo transition and how far out from the center can we trace the halo?**
  - ⦿ **Does the disk metallicity gradient change slope at large galactocentric radii?**
  - ⦿ **Do the identified tidal streams have clusters associated to them?**

**The combination of the number of filters, the sky coverage and the depth of this survey will make this project an unprecedented experiment for stellar population studies.**





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